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THE FIRST DETAILED ANALYSIS ON THE ASTROPHYSICAL PROPERTIES OF THE ECCENTRIC BINARY V990 HER

E. Kiran¹, V. Bakiş¹, H. Bakiş¹, and Ö. L. Değirmenci²

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ABSTRACT

We present accurate physical parameters of the eccentric binary system V990 Her which has an orbital period of $P = 8.193315 \pm 0.000003$ days using its photometric and spectroscopic data. The physical parameters of the components were derived as $T_{\text{eff1}} = 8000 \pm 200$ K, $T_{\text{eff2}} = 7570 \pm 200$ K, $M_1 = 2.01 \pm 0.07 M_{\odot}$, $M_2 = 1.83 \pm 0.03 M_{\odot}$, $R_1 = 2.22 \pm 0.02 R_{\odot}$, $R_2 = 2.12 \pm 0.01 R_{\odot}$, $\log(L_1/L_{\odot}) = 1.25 \pm 0.04$, $\log(L_2/L_{\odot}) = 1.12 \pm 0.05$. Our findings revealed that both components are slightly evolved from the zero-age main sequence with an age of 6.3×10^8 years. We estimated an apsidal motion with a period of $U = 14683 \pm 2936$ years in the system and the internal structure constants of the components were derived for the first time.

RESUMEN

Presentamos parámetros físicos precisos para la binaria excéntrica V990 Her, cuyo período es $P = 8.193315 \pm 0.000003$ días, mediante datos fotométricos y espectroscópicos. Para los parámetros físicos de las componentes obtuvimos los siguientes valores: $T_{\text{eff1}} = 8000 \pm 200$ K, $T_{\text{eff2}} = 7570 \pm 200$ K, $M_1 = 2.01 \pm 0.07 M_{\odot}$, $M_2 = 1.83 \pm 0.03 M_{\odot}$, $R_1 = 2.22 \pm 0.02 R_{\odot}$, $R_2 = 2.12 \pm 0.01 R_{\odot}$, $\log(L_1/L_{\odot}) = 1.25 \pm 0.04$, $\log(L_2/L_{\odot}) = 1.12 \pm 0.05$. Nuestras soluciones revelan que ambas componentes han evolucionado de la secuencia principal de edad cero, y que tienen edades de 6.3×10^8 años. Estimamos que el movimiento apsidal del sistema tiene un período de $U = 14683 \pm 2936$ y obtenemos por primera vez las constantes de la estructura interna de las componentes.

Key Words: binaries: eclipsing — stars: individual: V990 Her — techniques: photometric — techniques: spectroscopic

1. INTRODUCTION

Apsidal motion in detached eclipsing binaries with eccentric orbits is one of the methods used to examine the effects of the tidal distortion on the dynamical evolution of the binary and to test stellar evolution models. Orbital elements and absolute parameters of the components of such systems are required to estimate the internal structure constants (ISCs) of the components which are tests of stellar evolutionary models. Based on this motivation, we have examined the V990 Her binary system, which has not been studied in detail before.

V990 Her (HD169888, GSC 01581 01786) ($\alpha_{2000} = 18^{\text{h}}25^{\text{m}}56^{\text{s}}.71$, $\delta_{2000} = 21^{\circ}36'21''.34$) is an early type

(A0, Cannon & Pickering, 1993) eclipsing binary system, which was classified as an Algol type by Kazarovets et al. (1999). Later the system was listed in the eccentric eclipsing binary systems catalogue published by Otero et al. (2006) reporting the light elements as $T_0(HJD) = 2448048.755$, $P = 8^{\text{d}}.19329$. McDonald et al. (2012) reported the effective temperature $T_{\text{eff}} = 7514$ K and luminosity $L = 27.45 L_{\odot}$ of the system. The binary type of the system is given as detached main sequence system (DM) by Avvakumova et al. (2013).

In the following section (§2), we present information about the photometric and spectroscopic observations and reduction procedures of the data used in our analysis of the V990 Her system. The analysis of the data is given in § 3. In § 4 and § 5 we present the results and the discussion.

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2. OBSERVATIONS AND DATA REDUCTION

2.1. Photometric Data

In this study, both new observational and survey data (available to the public) were used for the photometric analysis.

Time series photometric observations of V990 Her were made between 2018-2019 using the 0.6 m automated telescope (T60) of the TÜBİTAK National Observatory (TUG), which is equipped with a 2048×2048 pixels FLI Pro-Line 3041-UV CCD camera with a pixel size of $15 \times 15 \mu\text{m}^2$. The Akdeniz University 0.25 m telescope (AUT25) which is equipped with a 2184×1472 pixels QSI632ws CCD camera with a pixel size of 6.8μ was also used. Magnitudes of the system in different passbands were extracted from the reduced **Bessel VRI** photometric frames with the **PHOT** package in IRAF³.

Additional photometric data of the system have been collected from four sources, the Hipparcos and Tycho Catalogues (Esa 1997), the All Sky Automated Survey (ASAS-3) (Pojmanski & Maciejewski 2004) Database, the Multi-site All-Sky CAM-era (MASCARA) (Talens et al. 2017) Survey and the QuickLook Pipeline (QLP) light curves (LCs) observed by the Transiting Exoplanet Survey Satellite (TESS) and produced by the MIT QuickLook Pipeline (QLP) (Huang et al. 2020). To obtain the times of minima to be used in the $O - C$ analysis in § 3.1 we used the MASCARA LCs published by Burggraaff et al. (2018) and the TESS QLP LCs (Sector 26) which are available at the Mikulski Archive for Space Telescopes (MAST) as a High Level Science Product. The photometric data taken from Hipparcos photometry and ASAS-3 were not sufficient to specify the minimum depth of the system. Therefore, we prefer to use only our T60 and AUT25 photometric data for the LCs analysis (see § 3.3 and § 3.6).

2.2. Spectroscopic Data

Spectroscopic data were collected by using the eShel spectrograph, which has a resolving power of 12000 covering the 4045-8100 Å wavelength range in 27 orders that is connected to the 60 cm telescope (UBT60) of Akdeniz University with a 20-m fiber. More information on the UBT60 telescope system is given by Bakış et al. (2020). We collected 29 spectra of V990 Her on different observing nights during

June-September 2017 and August-October 2018 seasons (see Table 1). In Table 1, the columns show the order of the observing night, mid-observing time of the spectra in heliocentric Julian date, total exposure time, S/N ratio value around 5500 Å and the orbital phase calculated with the ephemeris of the primary minimum given in equation (1).

IRAF was used for the reduction of the spectroscopic data and for the measurement of the components' radial velocities (RVs). In each observing night, bias, dark frames and flat spectra from a tungsten lamp were obtained for image reduction. Order aperture extractions and elimination of scattered light through the orders were made by using the IRAF ÉCHELLE package. The wavelength calibration was performed with thorium-argon lamp spectra which were taken before and after each object spectrum.

3. ANALYSIS

3.1. Improved Ephemerides and Apsidal Motion

Before attempting the photometric and spectroscopic analysis, the available ephemerides of the system given by Otero et al. (2006) were improved. Some photometric minimum data of V990 Her were published by Kreiner (2004), which were used, together with the newly obtained minima times, for the $O - C$ analysis. We measured two secondary and one primary times of minima from the MASCARA LCs and two primary and one secondary times of minima from the TESS QLP LCs. The measured original times of minima in units of BJD from TESS LCs were converted to HJD unit by using online applets⁴ (Eastman et al. 2010). Additionally, we obtained three primary and one secondary times of minima during our photometric observations. The Kwee & van Woerden (1956) method was used to determine the new times of minima and their corresponding errors for our observations. The obtained times of minima are listed in Table 2 in line with the literature; their uncertainty, epoch and $O - C$ values computed with the linear light elements in equation (1), type of minima ("pri" for primary and "sec" for secondary minimum), observational method ("ccd" for charge-coupled device observations) are also listed; the references are given in the last column.

When we performed a linear $O - C$ fit to all the primary and secondary minimum times given in Table 2 by using the unweighted least squares method, we found that the $O - C$ differences between the observed and the calculated times of minima regarding to the primary minimum behave in opposite phases

³IRAF is distributed by the National Optical Observatories, operated by the Association of Universities for research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

⁴<https://astrutils.astronomy.osu.edu/time/index.html>.

TABLE 1
JOURNAL OF SPECTROSCOPIC OBSERVATIONS*

No	Date (YYYYMMDD)	Julian Date (HJD)	Exp. Time (s)	<i>S/N</i>	Orbital Phase
1	20170516	2457890.3826	3000	40	0.178
2	20170606	2457911.4226	3000	40	0.746
3	20170611	2457916.4894	3000	110	0.365
4	20170612	2457917.5082	3000	110	0.489
5	20170622	2457927.4824	3000	40	0.706
6	20170702	2457937.4004	3000	110	0.917
7	20170703	2457938.4406	3000	50	0.044
8	20170704	2457939.5348	3000	100	0.177
9	20170709	2457944.4475	3000	100	0.777
10	20170711	2457946.4449	3000	135	0.021
11	20170719	2457954.4968	3000	100	0.003
12	20170720	2457955.3306	3000	110	0.105
13	20170723	2457958.4583	4200	140	0.487
14	20170802	2457968.3904	3000	140	0.699
15	20170918	2458015.3474	6600	50	0.430
16	20170920	2458017.3004	6000	50	0.669
17	20170928	2458025.2668	7200	40	0.641
18	20171011	2458038.2554	6000	50	0.226
19	20180828	2458359.3891	3000	110	0.421
20	20180829	2458360.3098	3000	110	0.533
21	20180913	2458375.3204	3000	70	0.365
22	20180917	2458379.3137	3000	80	0.853
23	20180918	2458380.3664	3000	110	0.981
24	20180919	2458381.3540	3000	100	0.102
25	20180922	2458384.2969	3000	110	0.461
26	20181002	2458394.3723	3000	40	0.691
27	20181003	2458395.3006	3000	40	0.804
28	20181005	2458397.2865	3000	140	0.046
29	20181007	2458399.2858	3000	125	0.290

*S/N ratio refers to the wavelength region around 5500 Å.

compared to those of the secondary minima. Then we applied the unweighted least squares fits to the primary and secondary eclipsing times separately and obtained the new linear ephemerides as follows (numbers in parentheses indicate errors in the last digit, and E is the number of cycles):

$$MinI(HJD) = 2448048.752(7) + 8.193314(6) \times E, \quad (1)$$

$$MinII(HJD) = 2448052.398(7) + 8.193316(7) \times E. \quad (2)$$

This result is a clear sign that the system shows apsidal motion. Therefore, we decided to perform an apsidal motion analysis using the times of minima (except for 2453177.7065 due to its large deviation from the expected location in the $O - C$ diagram) given in Table 2. To model the observed $O - Cs$ we used the equation given by Giménez & Bastero (1995) (equation 15 in their paper). We applied a differential correction method to represent the cor-

rections to the input values of the parameters, which are selected as adjustable, and used the unweighted least squares method to obtain them. Unfortunately, when we selected all five parameters as adjustable, we were unable to get a convergent solution. The main reason for this may be that the range covered by the times of minima available is short compared to the apsidal period of the system.

In order to reduce the number of the adjustable parameters and thus obtain an acceptable solution, we decided to determine the orbital eccentricity from the shape of the system's LC. In a LC of an eccentric eclipsing binary system, $e \cos w$ and $e \sin w$, the combinations of orbital eccentricity (e) with the longitude (w) of the periastron, are related with the amount of displacement of the secondary minimum with respect to the primary minimum and with the durations of the minima, respectively.

TABLE 2
TIMES OF MINIMA OF V990 HER

JD (Hel.) 2400000+	Error	E	$O - C$	Type	Method	Reference
48048.7550		0.0	0.0030	pri	ccd	Otero et al. (2006)
48052.401		0.5000	-0.4477	sec	ccd	Otero et al. (2006)
53177.7065*	0.0016	626.0	-0.0601	pri	ccd	Zakrzewski,B. (SuperWASP) unpublished
53746.7457	0.0024	695.5	-0.4477	sec	ccd	Zakrzewski,B. (ASAS-3) unpublished
53751.2885	0.0027	696.0	-0.0100	pri	ccd	Zakrzewski,B. (ASAS-3)unpublished
57192.4960	0.001	1116.0	0.0056	pri	ccd	This study (MASCARA)
57212.5300	0.002	1118.5	-0.4437	sec	ccd	This study (MASCARA)
57253.4970	0.001	1123.5	-0.4433	sec	ccd	This study (MASCARA)
58249.4300	0.010	1245.0	0.0021	pri	ccd	This study (AUT25)
58290.4000	0.01	1250.0	0.0055	pri	ccd	This study (AUT25)
58634.5000	0.01	1292.0	-0.0137	pri	ccd	This study (AUT25)
58646.3500	0.01	1293.5	-0.4537	sec	ccd	This study (AUT25)
59011.4045	0.0005	1338.0	-0.0016	pri	ccd	This study (TESS)
59019.5977	0.0006	1339.0	-0.0017	pri	ccd	This study (TESS)
59023.2540	0.001	1339.5	-0.4421	sec	ccd	This study (TESS)

*Not used in the apsidal motion analysis.

For relatively small eccentricities, these relationships are given as $e \cos w = \pi \delta / (1 + \csc^2 i)$ and $e \sin w = (D_2 - D_1) / (D_2 + D_1)$, where δ is the amount of displacement of the secondary minimum with respect to the primary minimum, i is the orbital inclination and D_1 and D_2 are the duration of the primary and secondary minimum, respectively. As a result, from the LCs (see Figure 7) of the system we obtained the values $e \cos w = -0.0867$ and $e \sin w = -0.0899$ which yield the orbital eccentricity as $e = 0.125$.

We managed to obtain a solution by keeping this value of eccentricity constant. The results are given in Table 3 associated with their internal errors. In Table 3, the parameters T_0 , P , e , w_0 and \dot{w} denote the initial epoch, orbital period, eccentricity, longitude of the periastron point at the initial epoch, and apsidal motion rate, respectively. The observed $O - C$ variations (symbols) of V990 Her are shown in Figure 1 concerning the theoretical curves (solid lines) computed with the apsidal motion elements given in Table 3.

According to the results given in Table 3, V990 Her has an apsidal motion with a period of $U = 14683 \pm 2936$ years. Due to the relatively long apsidal period compared with the time span of about 30 years which is covered by the observations, the apsidal motion parameters have larger errors. Therefore, the results given in Table 3 should be considered as a preliminary ones. Thus, for V990 Her more ob-

TABLE 3
APSIDAL MOTION PARAMETERS OF
V990 HER

Parameter	Value	Error
T_0 (HJD)	2448048.513	0.004
P (days)	8.193315	0.000003
e	0.125	(fixed)
w_0 (deg)	226.2	0.8
\dot{w} (deg cycle $^{-1}$)	0.00055	0.00011

servations, which will be performed in the future, are essential.

3.2. Spectroscopic Orbital Parameters

MgII, SiII, FeI, FeII and HI lines can be clearly seen in the spectrum of V990 Her. RVs were obtained from the orders containing the relatively sharp lines of MgII (4481Å), FeI (4957Å), FeI (5616Å) and HI (6563Å). The studied regions are shown in Figure 2. From top to bottom along the y-axis in Figure 2, the normalized spectra are plotted in accordance with the decreasing HJD times given in Table 1.

Gaussian profiles were applied to the central parts of the lines to measure the RVs and then a barycentric correction was applied to correct the RVs. Measured RVs were listed in Appendices A,

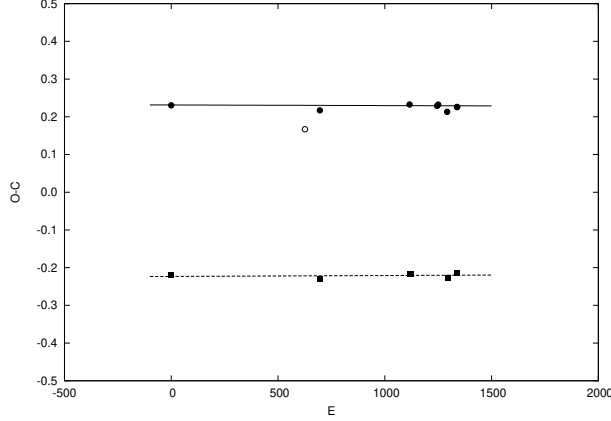


Fig. 1. The $O - C$ variations of V990 Her. Circles and squares represent primary and secondary minima, respectively. The open circle represents the minimum time which is not used in the analysis. Solid and dashed lines are for the theoretical model given in Table 3.

B, C and D. The least squares method has been applied to fit the RV equation to the measured RVs. We used the orbital period derived in equation (1) and fixed it during the iterations. The measured RVs and fitted curves (solid red lines) are shown in Figure 3 for each selected line separately. The spectroscopic orbital parameters, such as the longitude of the periastron, ω , orbital eccentricity, e , mass ratio, $q = K_1/K_2$ and velocity semi-amplitude of the components, $K_{1,2}$, were obtained from different lines and their arithmetic mean values are given in Table 4.

3.3. Photometric Elements

The previously mentioned three photometric data sets obtained in *VRI* bands were used for the photometric analysis of V990 Her to determine the photometric elements of the system and the light contributions of the components at different phases, which were used to model the component spectra in § 3.4 and § 3.5. The LC solutions were performed with the program PHOEBE 1.0 (Prša & Zwitter 2005), which is an extension of the WD program (Wilson & Devinney 1971).

The improved light elements given in equation (1) were used in the LCs analysis. The effective temperature of V990 Her was given as $T_{\text{eff}} = 7514$ K by McDonald et al. (2012). However, we obtained a more reliable temperature for the primary component as $T_{\text{eff1}} = 8000$ K directly from the spectrum of the system (see § 3.5). The temperature for the primary component T_{eff1} , the improved period given in § 3.1, the eccentricity e , and the mass

TABLE 4
PRELIMINARY ORBITAL PARAMETERS OF
V990 HER

Element	Value
ω (deg)	191.8 ± 0.7
e	0.09 ± 0.02
K_1 (km s $^{-1}$)	76.8 ± 0.1
K_2 (km s $^{-1}$)	87.1 ± 0.1
$q = K_1/K_2$	0.88 ± 0.01
V_γ (km s $^{-1}$)	-32.6 ± 3.5
$a \sin i$ (R_\odot)	26.4 ± 1.0

ratio q given in Table 4 were fixed during the LC solution. Convergence was allowed for the orbital inclination i , the longitude of the periastron point ω , the dimensionless surface potentials of the components Ω_1 and Ω_2 , the effective temperature of the secondary component $T_{\text{eff},2}$, and the relative monochromatic luminosity of the primary component l_1/l_{tot} . The bolometric albedos $A_{1,2}$ and gravitational darkening coefficients $g_{1,2}$ were adopted as 1.0, as suggested for early-type stars. The bolometric limb-darkening coefficients $x_{\text{bol},1,2}$ were computed automatically by PHOEBE 1.0 with the internal tables of the programme which were generated according to the models given by van Hamme (1993). The results of the solution are listed in Table 5.

3.4. Disentangling the Components Spectrum

One of the advantages of having numerous spectra taken at different phases is that they allow to obtain the spectrum of the components separately. Thus, we used the KOREL program provided by Hadrava (1995) which needs approximate orbital elements and the light contributions of the components at the phases in which the individual spectrum were taken. The orbital parameters and light contributions of the components obtained in § 3.2 and § 3.3 were used as initial values for the KOREL solution. The orbital period P given in equation (1) was used and fixed during the iterations. The KOREL solution gives the RVs (see Table 6) of the components and the orbital parameters (see Table 7) of the binary system. The KOREL solution agrees very well with those obtained from individual lines (see Table 4). The disentangled spectra in the H_α region for both components can be seen in Figure 4.

3.5. Atmosphere Modelling

We have a spectrum of the system in the middle of the secondary minimum of the LC where ≈ 45

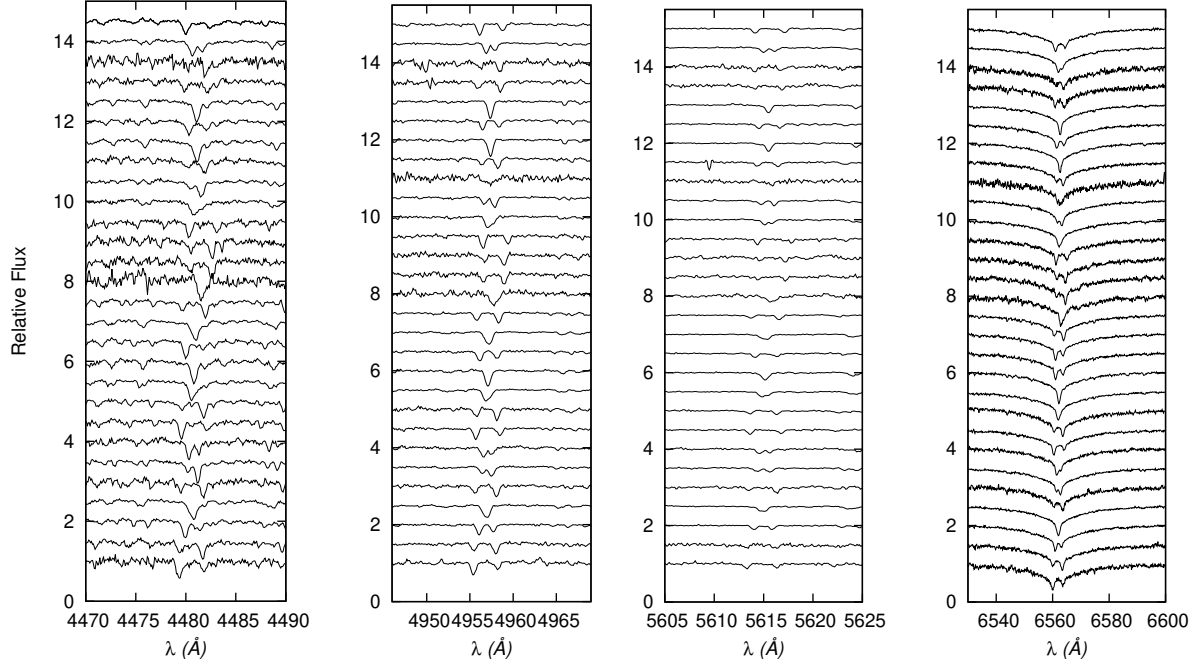


Fig. 2. Analyzed spectral lines. From top to bottom along the y-axis, the normalized spectra are plotted in accordance with the decreasing HJD times given in Table 1 in each panel.

TABLE 5
THE LIGHT CURVE MODEL OF V990 HER

Element	Primary	System	Secondary
P (d)		8.193314 (fixed)	
$PhaseShift$		-0.0278 ± 0.0001	
q		0.88 (fixed)	
e		0.09(fixed)	
w (rad)		3.50 ± 0.01	
i ($^{\circ}$)		85.48 ± 0.03	
T_{eff} (K)	8000 (fixed)		7544 ± 27
Ω	13.91 ± 0.08		11.92 ± 0.06
x_{bol}	0.198		0.199
y_{bol}	0.549		0.547
$l/l_{(1+2)}$ V band	0.526 ± 0.004		0.474 ± 0.004
$l/l_{(1+2)}$ R band	0.514 ± 0.004		0.486 ± 0.004
$l/l_{(1+2)}$ I band	0.504 ± 0.004		0.496 ± 0.004
r_{mean}	0.0772 ± 0.0004		0.0818 ± 0.0004
χ^2		0.777	

percent of the light of the secondary component is eclipsed. When the similar physical properties of the components (see Table 5) are taken into consideration, it can be regarded as a useful spectrum for a temperature estimation of the primary component.

We know that the components are not early type since there are no He I lines in the composite or dis-

entangled spectra, which means that their temperature is no higher than ≈ 10000 K, and therefore, the local thermodynamic equilibrium assumption could be used to construct the atmosphere models. ATLAS9 atmosphere models computed by Castelli & Kurucz (2004) and SYNTH program of Kurucz (1993) were used for the generation of the atmospheric models

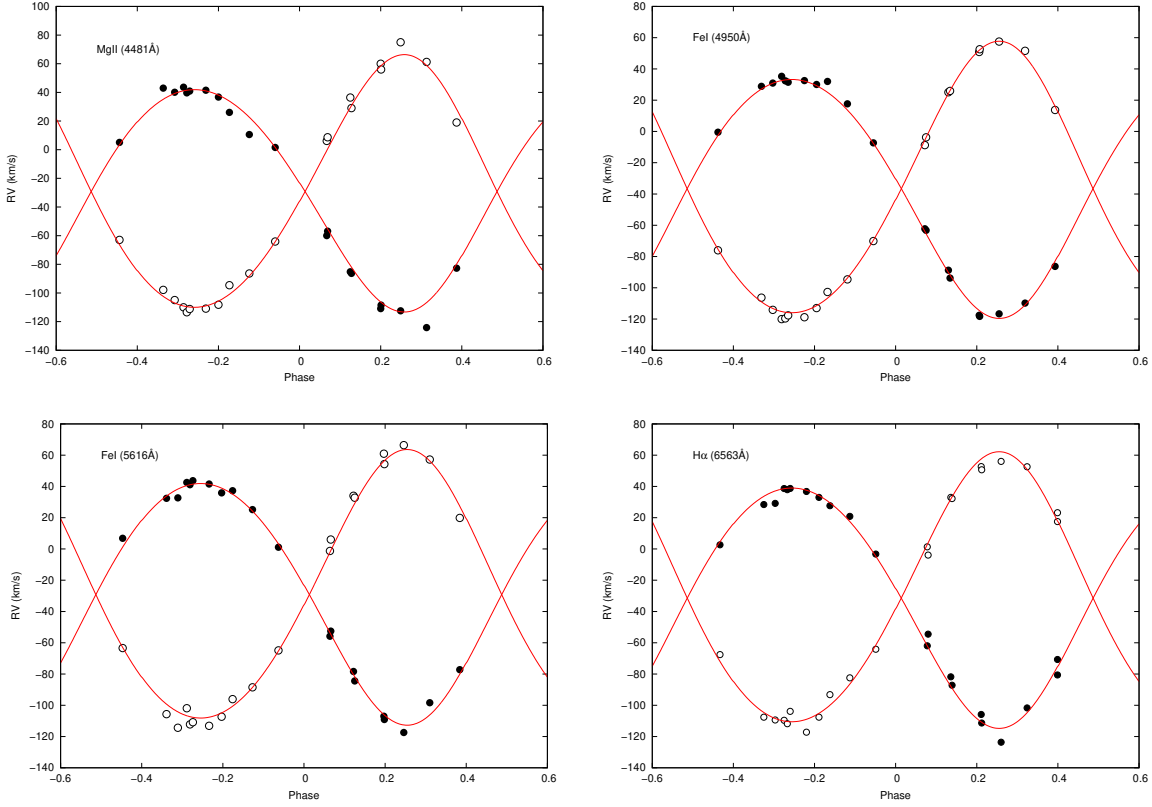


Fig. 3. RVs observations (symbols) and fitted models (red lines) for different spectral lines. Black dots and circles denote the primary and secondary components. The color figure can be viewed online.

and the synthetic spectra for the components. In the ATLAS9 code, $Z=0.03$ and $\xi = 2 \text{ km s}^{-1}$ were assumed for the metallicity and micro turbulence velocity, respectively. This method has been applied to the wavelength regions where the dominant lines reside. In Figure 5, the observational data and the synthetic spectra are given for four different wavelength regions as examples.

The atmospheric model parameters of the primary component, such as the temperature, $T_{\text{eff1}}=8000 \pm 200 \text{ K}$, the surface gravity, $\log g_1 = 4.0 \pm 0.1$ and the rotational velocity $v_{\text{rot1}}=28 \pm 5 \text{ km s}^{-1}$ were determined.

Moreover, we have a spectrum at orbital phase $\phi = 0.0$. Using a similar approach, the atmospheric model parameters of the secondary component were estimated as $T_{\text{eff2}}=7500 \pm 200 \text{ K}$, $\log g_2=4.0 \pm 0.1$ and $v_{\text{rot2}}=25 \pm 5 \text{ km s}^{-1}$. At three different orbital phases ($\phi = 0.699, 0.533, 0.290$), the synthetic spectra of the components were combined using their light contributions and their specific RVs. The observed spec-

tra taken at the indicated maximum phases and the combined theoretical spectra are shown together in Figure 6.

3.6. Simultaneous Radial Velocity and Light Curves Solution

The RVs of the components obtained from the KOREL solution and the multi-colour LCs of the system were analyzed simultaneously with PHOEBE 1.0. The initial orbital parameters were used from the averaged values given in Tables 4 and 7. We used the improved light elements given in equation (1). The orbital period P and the effective temperature of the primary component T_{eff1} were fixed as in § 3.3, the other orbital and photometric parameters were all adjusted in the analyses.

The results of the simultaneous solution are listed in Table 8. The observed LCs in *VRI* bands and the radial velocity curves (RVCs) of the components are shown along with the theoretical models in Figure 7. The Roche geometry of V990 Her at phases $\phi=0.0, 0.25$ and 0.45 can be seen in Figure 8.

TABLE 6
ADOPTED RVs OF THE COMPONENTS OF V990 HER

No	HJD (+2400000)	Phase	Primary RV km s^{-1}	$(O - C)_1$ km s^{-1}	Secondary RV km s^{-1}	$(O - C)_2$ km s^{-1}
1	57890.3826	0.178	-113.42	-4.06	51.26	-5.99
2	57911.4226	0.746	32.72	-5.56	-119.83	-9.04
3	57916.4894	0.365	-84.21	-4.91	14.56	-8.48
4	57917.5082	0.489	-	-	-	-
5	57927.4824	0.706	38.85	-0.81	-112.49	-0.13
6	57937.4004	0.917	-2.22	1.07	-65.03	-1.55
7	57938.4406	0.044	-62.92	-3.32	1.15	0.55
8	57939.5348	0.177	-111.31	-2.14	51.65	-5.38
9	57944.4475	0.777	38.67	3.79	-107.78	-0.85
10	57946.4449	0.021	-	-	-	-
11	57954.4968	0.003	-	-	-	-
12	57955.3306	0.105	-85.53	1.04	32.20	0.88
13	57958.4583	0.487	-	-	-	-
14	57968.3904	0.699	43.63	4.10	-107.86	4.35
15	58015.3474	0.430	-	-	-	-
16	58017.3004	0.669	63.50	25.82	-83.33	26.77
17	58025.2668	0.641	57.11	22.97	-103.97	2.10
18	58038.2554	0.226	-77.16	37.56	86.45	23.10
19	58359.3891	0.421	-	-	-	-
20	58360.3098	0.533	2.17	-1.44	-71.89	-0.57
21	58375.3204	0.365	-	-	-	-
22	58379.3137	0.853	20.40	2.04	-89.44	-1.31
23	58380.3664	0.981	-	-	-	-
24	58381.3540	0.102	-83.37	1.92	32.56	2.71
25	58384.2969	0.461	-	-	-	-
26	58394.3723	0.691	39.51	0.27	-110.91	0.97
27	58395.3006	0.804	29.90	-0.34	-100.61	1.02
28	58397.2865	0.046	-58.81	2.11	8.22	6.12
29	58399.2858	0.290	-113.89	-6.95	45.34	-9.15

4. RESULTS

4.1. Physical Properties and Evolutionary Status

TABLE 7

ORBITAL PARAMETERS OF V990 HER
DERIVED FROM KOREL SOLUTION

Element	Value
T_{per} (days)	2457883.010 ± 0.003
ω (deg)	200.7 ± 0.3
e	0.085 ± 0.008
K_1 (km s^{-1})	77.2 ± 0.4
K_2 (km s^{-1})	87.9 ± 0.6
$q = K_1/K_2$	0.880 ± 0.002

Using the parameters given in Tables 7 and 8, the astrophysical parameters of the components and the distance of V990 Her were derived and are given in Table 9. The masses of the components ($2.01 \pm 0.07 M_{\odot}$ and $1.83 \pm 0.03 M_{\odot}$) correspond to main sequence stars of spectral type between A3-A5 (Straizys & Kuriliene 1981). The evolution status of the components according to the T_{eff} and $\log L$ values given in Table 9 is shown in Figure 9 along with the theoretical evolutionary tracks taken from Bressan et al. (2012). The isochrone in Figure 9 was taken from Girardi et al. (2000).

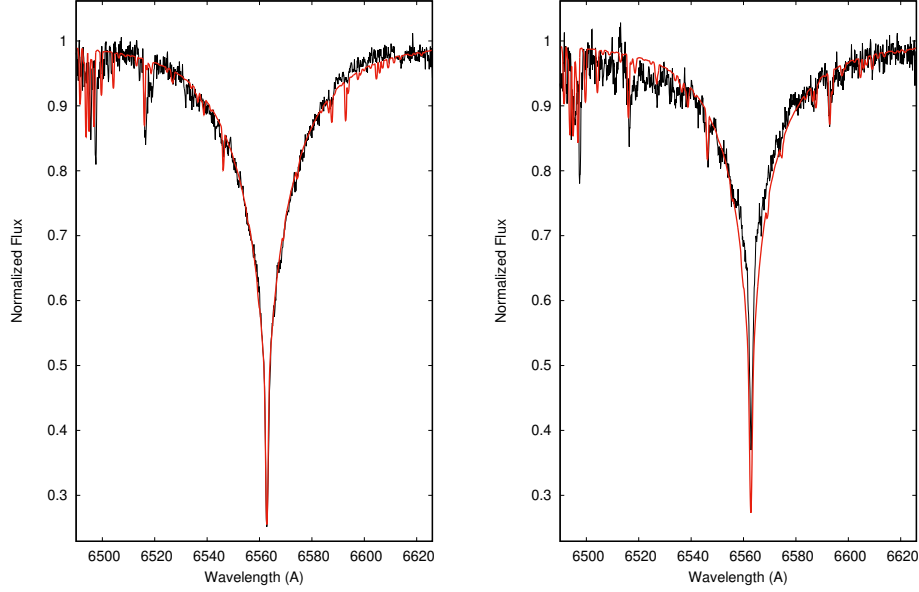


Fig. 4. Disentangled spectra in H_α region for the primary (left panel) and the secondary (right panel) components of V990 Her obtained from KOREL solution. The red solid lines represent the synthetic spectra calculated in § 3.5. The color figure can be viewed online.

TABLE 8
THE SIMULTANEOUS SOLUTION OF RVCS AND LCS OF V990 HER

Element	Primary	System	Secondary
P (d)		8.193314 (fixed)	
$PhaseShift$		-0.0278 ± 0.0001	
a (R_\odot)		26.8 ± 0.2	
q		0.912 ± 0.005	
V_γ (km s^{-1})		-31.8 ± 0.3	
e		0.089 ± 0.001	
w ($^\circ$)		198 ± 1	
i ($^\circ$)		85.46 ± 0.03	
T_{eff} (K)	8000 (fixed)		7570 ± 31
Ω	13.06 ± 0.07		12.72 ± 0.05
x_{bol}	0.198		0.199
y_{bol}	0.549		0.547
$l/l_{(1+2)}$ V band	0.575 ± 0.006		0.425
$l/l_{(1+2)}$ R band	0.564 ± 0.005		0.436
$l/l_{(1+2)}$ I band	0.554 ± 0.006		0.447
r_{mean}	0.0828 ± 0.0005		0.0792 ± 0.0004
χ^2		0.800	

Bolometric corrections (BC) for the primary and secondary components were calculated as 0.027 and 0.034 mag using $T_{\text{eff}}-BC$ tables of Flower (1996). The bolometric magnitudes M_{bol} given in Table 9 and BC s were used to derive the absolute magni-

tudes of the components as $M_{V1} = 1^m.60 \pm 0.11$ and $M_{V2} = 1^m.93 \pm 0.12$, respectively. By using the $V_J = V_T - 0.09 \times (B_T - V_T)$ relation and Tycho magnitudes ($B_T = 7^m.950$, $V_T = 7^m.710$) given by Esa (1997), we derived the total visual magnitude of the

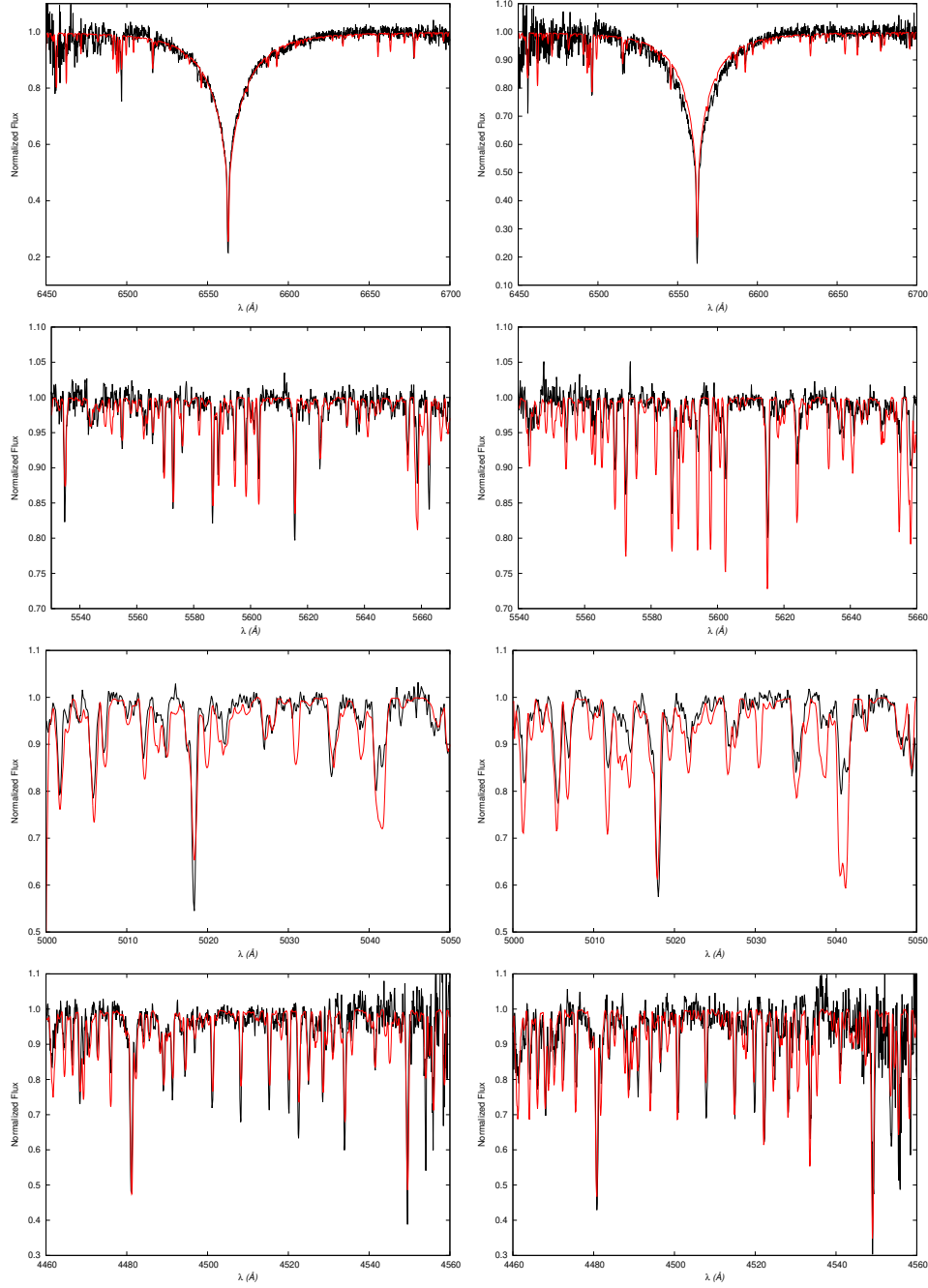


Fig. 5. The observed spectra taken during the secondary (left panel) and primary (right panel) minimum. The synthetic models are shown with red solid lines. The color figure can be viewed online.

system as $m_V = 7^m.688$. To obtain the $E(B - V)$ color excess and the distance of the system we used the total $u'g'r'i'z'$ magnitudes of the system given by Pickles & Depagne (2010) in the following transformation relationships of Smith et al. (2002):

$$V = g' - 0.55(g' - r') - 0.03, \quad (3)$$

$$(V - R_C) = 0.59(g' - r') + (0.11), \quad (4)$$

$$(R - I)_C = 1.00(r' - i') + (0.21) \text{ for } (r' - i') < 0.95. \quad (5)$$

Using the light contribution of the components given in Table 8 together with the I_C magnitudes of the system estimated by equations 3-5 and the

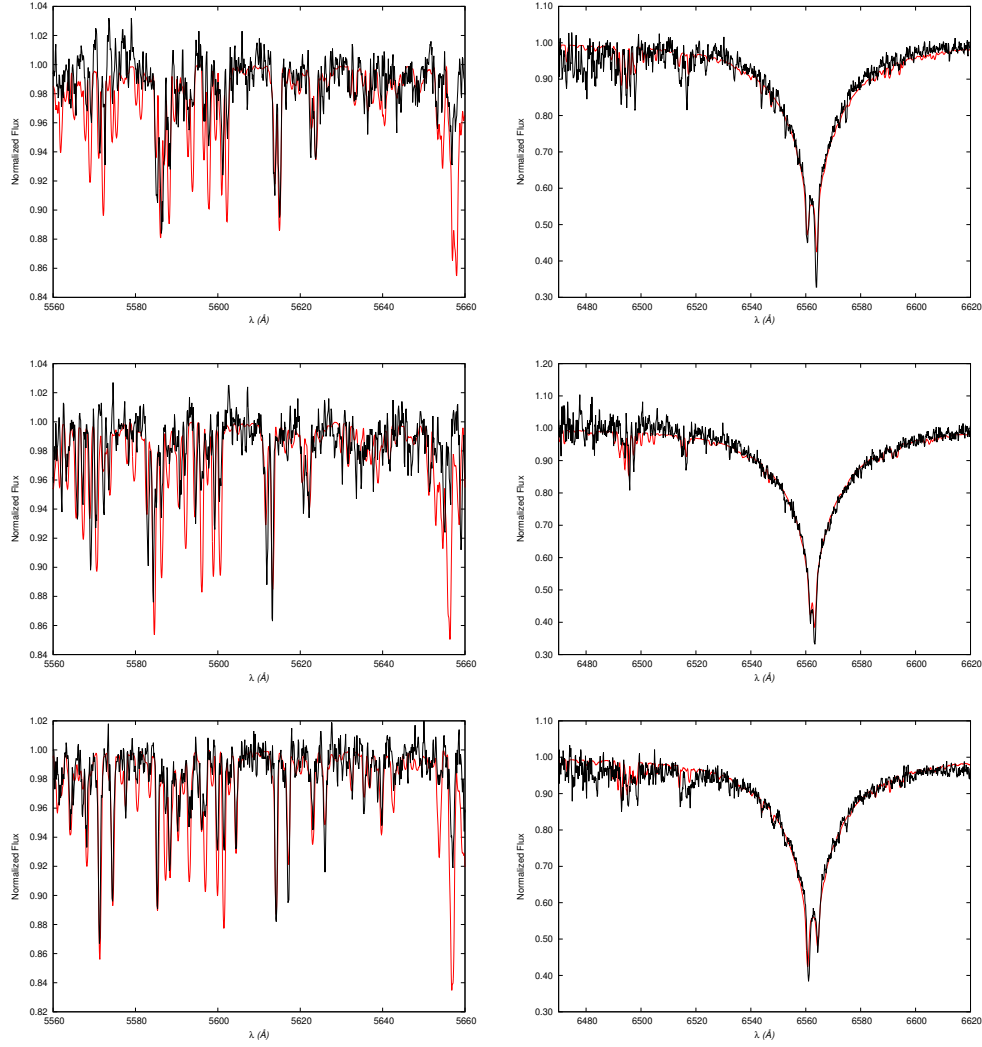


Fig. 6. The composite synthetic models (red solid lines) fitted to the observed spectra around Fe I (5616Å) and H I (6563Å) lines. Each row refers to orbital phases at $\phi = 0.699$ (upper panels), 0.533 (middle panels) and 0.290 (bottom panels). The color figure can be viewed online.

$(B - V) - (V - I_C)$ tables given by Cousins (1981), we obtained the $(B - V)$ colors as 0.185 and 0.250 mag for the primary and secondary components, respectively. For the unreddened $(B - V)_0$ colors of the components we used the $T_{\text{eff}} - (B - V)_0$ tables of Flower (1996) and obtained the values of $(B - V)_0$ as 0.162 and 0.234 mag, which correspond to $E(B - V)$ values of 0.023 and 0.017 mag for the primary and secondary components, respectively. Thus, we adopted an average color excess of $E(B - V) = 0.020$ mag and as a result we obtained the distance of the system as $d = 212 \pm 21$ pc. Alternatively, an $E(B - V) = 0^m.066$ value estimated from stislism maps given by Kervella et al. (2019) corresponds to a distance of $d = 198 \pm 19$ pc. However,

our photometric distance value matches better the value of 235 pc given by Gaia which corresponds to a parallax of $\mu = 4.25 \pm 0.08$ mas (Gaia Collaboration et al. 2016; Lindegren et al. 2021) within the given uncertainty box.

4.2. Apsidal Motion and Internal Structure Constant

To test the observed preliminary results concerning the apsidal motion of V990 Her given in § 3.1, we can use the results of previous section. As is known, the total apsidal motion in a binary system occurs mainly for two contributions: (1) a classical part which is due to tidal interactions of the oblate components, and (2) a relativistic contribution which is a direct result of the General Relativity Theory.

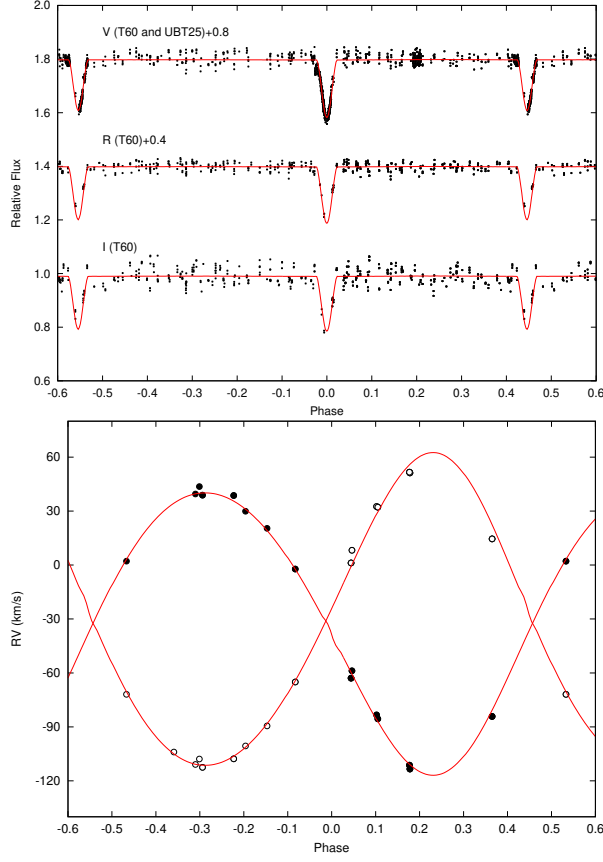


Fig. 7. The theoretical representations of the observed LCs of the system in VRI bands (upper panel) and the observed RVCs of the components (lower panel). The V and R band LCs are shifted 0.8 and 0.4 mag, respectively, for a better illustration. The color figure can be viewed online.

TABLE 9

THE BASIC PROPERTIES OF V990 HER

Element	Unit	Primary	Secondary
M	M_{\odot}	2.01 ± 0.07	1.83 ± 0.03
R	R_{\odot}	2.22 ± 0.02	2.12 ± 0.01
T_{eff}	K	8000 ± 200	7570 ± 200
$\log L$	L_{\odot}	1.25 ± 0.04	1.12 ± 0.05
M_{bol}	mag	1.6 ± 0.1	2.0 ± 0.1
$\log g$	cgs	4.06 ± 0.01	4.04 ± 0.01
V_{rot}	km s^{-1}	28 ± 5	25 ± 5
V_{syn}	km s^{-1}	13.6 ± 0.1	13.1 ± 0.1
d	pc	212 ± 21	

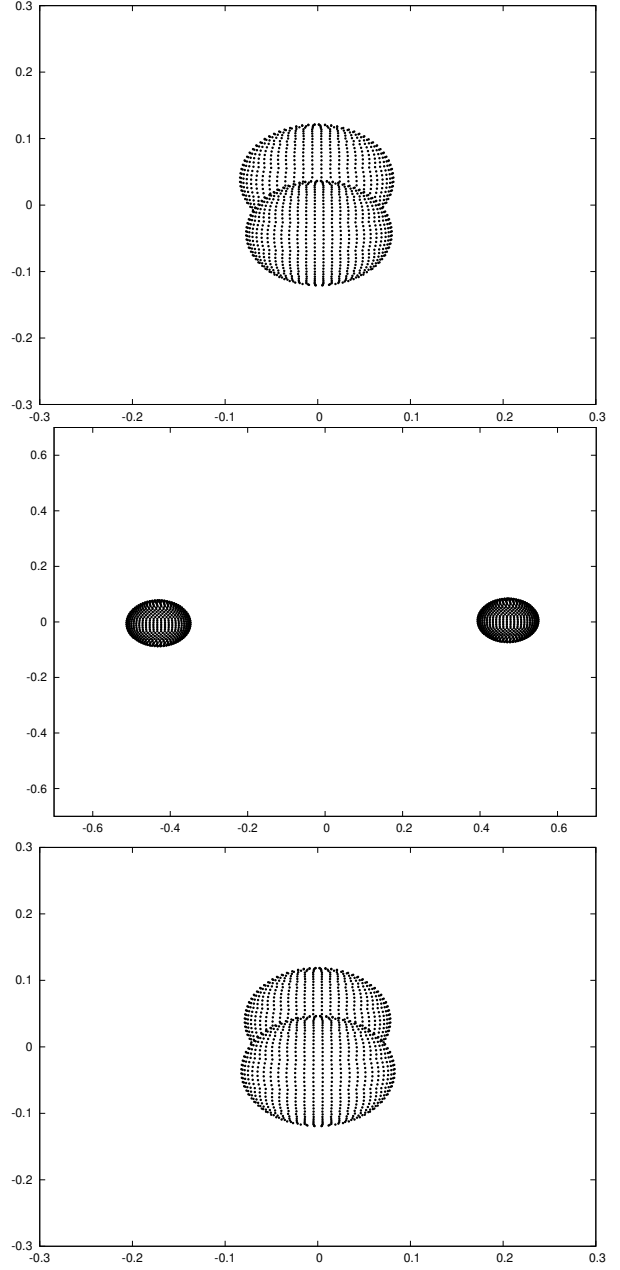


Fig. 8. The Roche geometry of V990 Her at the phases $\phi = 0.00$ (upper panel), $\phi = 0.25$ (middle panel) and $\phi = 0.45$ (lower panel).

The relativistic contribution \dot{w}_{rel} (in degrees per orbital cycle) to total apsidal motion is given by (see Giménez 1985):

$$\dot{w}_{\text{rel}} = 5.45 \times 10^{-4} \frac{1}{1 - e^2} \left(\frac{m_1 + m_2}{P} \right)^{2/3}, \quad (6)$$

which was first derived by Levi-Civita (1937). Here, $m_{1,2}$ are the masses of the components in units of so-

lar mass, e is the orbital eccentricity and P is the orbital period in units of days. By using the related parameters given in Tables 3 and 9 we obtained $\dot{w}_{rel} = 0^\circ.00033 \text{ cycle}^{-1}$ and so $\dot{w}_{cl} = 0^\circ.00022 \text{ cycle}^{-1}$.

On the other hand, the relation for the classical part of the apsidal motion rate was given by Kopal (1978) as follows

$$\frac{\dot{w}_{cl}}{360} = c_{21}k_{21} + c_{22}k_{22}, \quad (7)$$

where k_{2i} ($i=1, 2$) are internal structure constants (ISCs) of the components, while c_{2i} are as follows

$$c_{2i} = \left[\left(\frac{w_i}{w_k} \right)^2 \left(1 + \frac{m_{3-i}}{m_i} \right) f(e) + 15 \frac{m_{3-i}}{m_i} g(e) \right] (R_i/a)^5. \quad (8)$$

Here R_i/a and w_i/w_k are the relative radii and rotational angular velocity for the corresponding component, respectively, $w_k = \frac{2\pi}{P}$ is the orbital (Keplerian) angular velocity and P is the orbital period of the system. The f and g functions of the orbital eccentricity e are given by

$$f(e) = (1 - e^2)^{-2}, \quad (9)$$

$$g(e) = (1 - e^2)^{-5} (1 + 1.5e^2 + 0.125e^4). \quad (10)$$

Using the apsidal and absolute parameters given in Tables 3 and 9 we obtained the observed average ISC for the system as $\bar{k}_2 = 0.0035 \pm 0.0008$.

To obtain the theoretical ISCs of the components we used the stellar models given by Claret & Gimenez (1992), which assume $X=0.649$, $Z=0.03$ and take mass loss and core overshooting into account. By the interpolation we obtained $\log L=1.30$ and $1.13L_\odot$, $T_{\text{eff}}=8222$ and $7621K$, $\text{Age} = 5.72 \times 10^8$ and 7.89×10^8 years and $k_2 = 0.0032$ and 0.0031 for the primary and secondary components, respectively. These results correspond to a value of $\bar{k}_2(\text{theo}) = 0.0032$ and $\dot{w}_{cl}(\text{theo}) = 0^\circ.00019 \text{ cycle}^{-1}$. Consequently, the theoretical value of the total apsidal motion rate is obtained as $\dot{w}(\text{theo}) = 0^\circ.00052 \text{ cycle}^{-1}$. On the other hand, by taking the weighted average of the ages of the components with the corresponding c_i we obtained the age of the system as 6.74×10^8 years.

5. DISCUSSION

The absolute parameters of the eccentric binary V990 Her has been obtained with a good precision; ≤ 1 per cent for masses and radii. This allowed us to

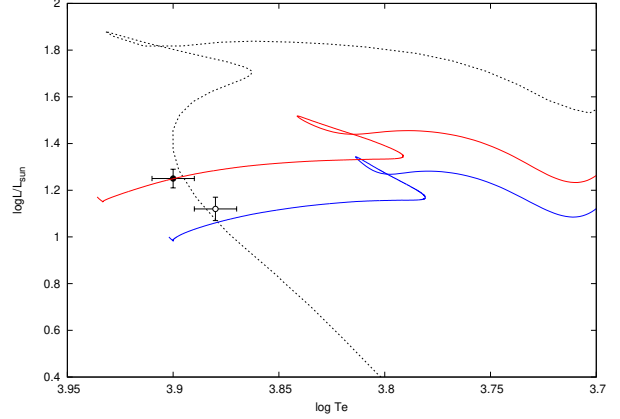


Fig. 9. Comparison of components with evolution models (Black dot and circle represent the primary and secondary components, red and blue solid lines indicates evolution models for $Z = 0.03$ for a $2.00M_\odot$ and $1.825M_\odot$ star individually). Black dashed line represents isochrone corresponds 6.3×10^8 yrs for $Z = 0.03$. The color figure can be viewed online.

see the evolutionary status of the system. The system consists of two early A-type main-sequence stars with an age of 6.3×10^8 years estimated by isochrone fitting, which is compatible with 6.74×10^8 years obtained from the apsidal motion analysis. The circularization and synchronization time scales were computed as 1.5×10^8 and 7.6×10^{10} years according to Zahn (Zahn 1977). The averaged age of the system is larger than the circularization and smaller than the synchronization time scales, which explains the fast rotation speed of the components.

Different $E(B - V)$ color excess values of $0^m.166$ (Jayasinghe et al. 2018), $0^m.11$ (Gontcharov & Mosenkov 2017) and $0^m.066$ (Kervella et al. 2019) are also available in the literature. According to these values, we estimated the distance of the system as 172 ± 17 , 186 ± 18 and 198 ± 19 pc, respectively. Our photometric distance ($d = 212 \pm 21$ pc), estimated in § 4.1, is closer to the one given by Gaia than the distances determined using the above color excesses.

The apsidal motion in V990 Her puts it into a special place and allows us to estimate the structure of the component stars. However, the time span of the available data we have is a very small portion ($\approx 0.2\%$) of the apsidal cycle. This situation creates problems in the apsidal motion analysis of the system and makes it difficult to determine the apsidal motion rate and orbital eccentricity together. For

now, based on our findings we can say that the orbital eccentricity of the system has a value between 0.09 and 0.13. More accurate results can be obtained with more observations to be made in the future. Lastly, the apsidal motion of V990 Her has a period of about 14700 years and is predominantly (60%) due to relativity.

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APPENDIX A. RADIAL MEASURED FROM MgII (4481.13Å)

No	Time (HJD) (+2400000)	Phase	RV_{pri} km s ⁻¹	eRV_{pri}	RV_{sec} km s ⁻¹	eRV_{sec}
1	57890.3826	0.17819	-108.49	1.26	56.00	1.80
2	57911.4226	0.74614	41.53	1.15	-110.89	0.39
3	57916.4894	0.36455	-82.72	10.86	18.99	2.88
4	57917.5082	0.48889	-21.62	2.19	-	-
5	57927.4824	0.70625	40.92	1.92	-111.20	0.62
6	57937.4004	0.91675	1.59	4.25	-64.10	4.02
7	57938.4406	0.04371	-59.96	0.08	6.12	0.19
8	57939.5348	0.17726	-110.89	0.41	60.00	3.53
9	57944.4475	0.77685	36.74	0.48	-108.18	0.52
10	57946.4449	0.02064	-	-	-44.75	2.20
11	57954.4968	0.00338	-	-	-31.85	0.50
12	57955.3306	0.10514	-86.31	0.23	29.01	1.44
13	57958.4583	0.48688	-22.08	2.41	-	-
14	57968.3904	0.69910	39.67	0.34	-113.48	0.19
15	58015.3474	0.43024	-30.26	1.67	-	-
16	58017.3004	0.66860	40.08	4.93	-104.93	6.02
17	58025.2668	0.64091	42.99	3.21	-97.84	0.20
18	58038.2554	0.22618	-112.45	3.80	75.00	1.59
19	58359.3891	0.42077	-40.57	5.14	-	-
20	58360.3098	0.53315	5.07	0.10	-62.93	0.43
21	58375.3204	0.36521	-	-	-	-
22	58379.3137	0.85258	10.56	27.31	-86.40	1.34
23	58380.3664	0.98107	-	-	-25.82	0.56
24	58381.3540	0.10161	-85.17	11.40	36.38	0.91
25	58384.2969	0.46079	-27.83	0.35	-	-
26	58394.3723	0.69050	43.64	0.21	-110.00	1.16
27	58395.3006	0.80380	26.02	6.82	-94.60	10.76
28	58397.2865	0.04618	-56.89	0.24	8.67	2.10
29	58399.2858	0.29019	-124.13	38.11	61.24	1.43

APPENDIX B. RVs MEASURED FROM FeI (4957.59Å)

No	Time (HJD) (+2400000)	Phase	RV_{pri} km s ⁻¹	eRV_{pri}	RV_{sec} km s ⁻¹	eRV_{sec}
1	57890.3826	0.17819	-118.29	2.06	52.60	0.97
2	57911.4226	0.74614	32.55	0.79	-118.86	1.15
3	57916.4894	0.36455	-86.34	1.11	13.73	1.02
4	57917.5082	0.48889	-32.50	2.46	-	-
5	57927.4824	0.70625	31.50	0.23	-117.69	1.28
6	57937.4004	0.91675	-7.28	1.02	-70.12	0.32
7	57938.4406	0.04371	-62.26	1.16	-8.76	1.67
8	57939.5348	0.17726	-117.61	0.68	50.84	0.59
9	57944.4475	0.77685	30.04	0.55	-113.01	2.26
10	57946.4449	0.02064	-	-	-45.42	3.14
11	57954.4968	0.00338	-	-	-37.40	0.24
12	57955.3306	0.10514	-93.81	1.26	25.91	1.38
13	57958.4583	0.48688	-32.73	1.90	-	-
14	57968.3904	0.69910	32.42	1.00	-119.61	1.20
15	58015.3474	0.43024	-41.34	1.05	-	-
16	58017.3004	0.66860	30.95	1.38	-114.16	2.28
17	58025.2668	0.64091	28.89	2.63	-106.35	1.36
18	58038.2554	0.22618	-116.67	1.37	57.50	0.75
19	58359.3891	0.42077	-49.35	2.78	-	-
20	58360.3098	0.53315	-0.49	1.25	-76.03	1.63
21	58375.3204	0.36521	-	-	-	-
22	58379.3137	0.85258	17.70	0.34	-94.66	1.87
23	58380.3664	0.98107	-	-	-33.01	0.35
24	58381.3540	0.10161	-88.77	0.70	25.08	0.87
25	58384.2969	0.46079	-34.77	0.61	-	-
26	58394.3723	0.69050	35.27	0.33	-120.04	3.05
27	58395.3006	0.80380	31.99	1.37	-102.69	4.17
28	58397.2865	0.04618	-63.24	0.13	-3.75	0.58
29	58399.2858	0.29019	-109.79	0.34	51.62	0.75

APPENDIX C. RVs MEASURED FROM FeI (5615.64Å)

No	Time (HJD) (+2400000)	Phase	RV_{pri} km s ⁻¹	eRV_{pri}	RV_{sec} km s ⁻¹	eRV_{sec}
1	57890.3826	0.17819	-109.10	1.38	54.27	1.41
2	57911.4226	0.74614	41.54	2.41	-113.12	4.36
3	57916.4894	0.36455	-77.21	0.13	19.85	0.54
4	57917.5082	0.48889	-19.13	5.37	-	-
5	57927.4824	0.70625	43.71	2.79	-110.86	0.61
6	57937.4004	0.91675	1.13	1.58	-64.91	2.58
7	57938.4406	0.04371	-55.86	0.91	-1.26	0.36
8	57939.5348	0.17726	-107.00	0.05	60.96	1.23
9	57944.4475	0.77685	35.90	0.13	-107.32	1.57
10	57946.4449	0.02064	-	-	-20.40	0.49
11	57954.4968	0.00338	-	-	-31.26	1.15
12	57955.3306	0.10514	-84.44	1.80	32.79	2.35
13	57958.4583	0.48688	-28.89	1.10	-	-
14	57968.3904	0.69910	41.09	0.45	-112.25	0.05
15	58015.3474	0.43024	-44.84	3.02	-	-
16	58017.3004	0.66860	32.70	31.86	-114.36	0.98
17	58025.2668	0.64091	32.42	0.32	-105.70	1.92
18	58038.2554	0.22618	-117.41	0.94	66.41	0.73
19	58359.3891	0.42077	-43.17	0.16	-	-
20	58360.3098	0.53315	6.84	0.54	-63.41	0.27
21	58375.3204	0.36521	-	-	-	-
22	58379.3137	0.85258	25.18	0.29	-88.52	1.56
23	58380.3664	0.98107	-	-	-25.86	0.48
24	58381.3540	0.10161	-78.36	0.25	34.04	2.18
25	58384.2969	0.46079	-25.95	1.13	-	-
26	58394.3723	0.69050	42.53	6.13	-101.88	4.49
27	58395.3006	0.80380	37.28	24.73	-96.06	9.58
28	58397.2865	0.04618	-52.55	0.54	6.05	0.20
29	58399.2858	0.29019	-98.37	0.77	57.20	2.44

APPENDIX D. RVs MEASURED FROM HI (6562.79Å)

No	Time (HJD) (+2400000)	Phase	RV_{pri} km s ⁻¹	eRV_{pri}	RV_{sec} km s ⁻¹	eRV_{sec}
1	57890.3826	0.17819	-111.34	2.24	50.65	2.07
2	57911.4226	0.74614	36.79	1.16	-117.24	2.17
3	57916.4894	0.36455	-80.59	2.66	23.18	4.54
4	57917.5082	0.48889	-28.70	0.71	-	-
5	57927.4824	0.70625	38.73	2.35	-103.90	6.91
6	57937.4004	0.91675	-3.24	0.75	-64.17	1.30
7	57938.4406	0.04371	-61.95	1.64	1.37	0.68
8	57939.5348	0.17726	-105.89	8.05	52.67	4.79
9	57944.4475	0.77685	33.00	1.13	-107.61	1.60
10	57946.4449	0.02064	-	-	-39.90	1.13
11	57954.4968	0.00338	-	-	-31.83	0.44
12	57955.3306	0.10514	-87.17	0.12	32.26	0.88
13	57958.4583	0.48688	-27.62	1.97	-	-
14	57968.3904	0.69910	37.92	0.31	-111.79	1.05
15	58015.3474	0.43024	-47.93	3.27	-	-
16	58017.3004	0.66860	29.16	2.61	-109.43	11.90
17	58025.2668	0.64091	28.44	4.96	-107.64	1.94
18	58038.2554	0.22618	-123.64	0.12	56.01	2.88
19	58359.3891	0.42077	-38.60	2.26	-	-
20	58360.3098	0.53315	2.67	0.43	-67.58	1.20
21	58375.3204	0.36521	-70.70	20.16	17.51	31.75
22	58379.3137	0.85258	20.87	0.63	-82.44	0.28
23	58380.3664	0.98107	-	-	-29.16	0.33
24	58381.3540	0.10161	-81.79	1.36	33.02	0.30
25	58384.2969	0.46079	-27.06	0.96	-	-
26	58394.3723	0.69050	38.68	3.70	-109.63	1.22
27	58395.3006	0.80380	27.69	0.45	-93.17	2.33
28	58397.2865	0.04618	-54.51	0.25	-3.93	1.20
29	58399.2858	0.29019	-101.71	0.05	52.61	0.48

REFERENCES

- Avvakumova, E. A., Malkov, O. Y., & Kniazev, A. Y. 2013, *AN*, 334, 860
- Bakiş, V., Eker, Z., Sarı, O., Yücel, G., & Sonbaş, E. 2020, *MNRAS*, 496, 2605
- Bressan, A., Marigo, P., Girardi, L., et al. 2012, *MNRAS*, 427, 127
- Burggraaff, O., Talens, G. J. J., Spronck, J., et al. 2018, *A&A*, 617, 32
- Cannon, A. J. & Pickering, E. C. 1993, *VizieR Online Data Catalog: III/135A*
- Castelli, F. & Kurucz, R. L. 2004, *arXiv Astrophysics e-prints*, astro-ph/0405087
- Claret, A. & Gimenez, A. 1992, *A&AS*, 96, 255
- Cousins, A. W. J. 1981, *SAAOC*, 6, 4
- Eastman, J., Siverd, R., & Gaudi, B. S. 2010, *PASP*, 122, 935
- ESA. 1997, *The Hipparcos and Tycho Catalogues*, ESA SP-1200
- Flower, P. J. 1996, *ApJ*, 469, 355
- Gaia Collaboration, Prusti, T., de Bruijne, J. H. J., et al. 2016, *A&A*, 595, 1
- Giménez, A. 1985, *ApJ*, 297, 405
- Giménez, A. & Bastero, M. 1995, *Ap&SS*, 226, 99
- Girardi, L., Bressan, A., Bertelli, G., & Chiosi, C. 2000, *A&AS*, 141, 371
- Gontcharov, G. A. & Mosenkov, A. V. 2017, *MNRAS*, 472, 3805
- Hadrava, P. 1995, *A&AS*, 114, 393
- Huang, C. X., Vanderburg, A., Pál, A., et al. 2020, *RNAAS*, 4, 204
- Jayasinghe, T., Kochanek, C. S., Stanek, K. Z., et al. 2018, *MNRAS*, 477, 3145
- Kazarovets, E. V., Samus, N. N., Durlevich, O. V., et al. 1999, *IBVS*, 4659, 1
- Kervella, P., Arenou, F., Mignard, F., & Thévenin, F. 2019, *A&A*, 623, 72
- Kopal, Z. 1978, *Dynamics of close binary systems* (Dordrecht: Astrophysics and Space Science Library, Springer)
- Kreiner, J. M. 2004, *AcA*, 54, 207
- Kurucz, R. L. 1993, *SYNTHES Spectrum Synthesis Programs and Line Data* (Cambridge, MA: Smithsonian Astrophysical Observatory)
- Kwee, K. K. & van Woerden, H. 1956, *BAN*, 12, 327
- Levi-Civita, T. 1937, *Am. J. Math.*, 59, 255
- Lindgren, L., Klioner, S. A., Hernández, J., et al. 2021, *A&A*, 649, A2
- McDonald, I., Zijlstra, A. A., & Boyer, M. L. 2012, *MNRAS*, 427, 343
- Otero, S. A., Wils, P., Hoogeveen, G., & Dubovsky, P. A. 2006, *IBVS*, 5681, 1
- Pickles, A. & Depagne, É. 2010, *PASP*, 122, 1437
- Pojmanski, G. & Maciejewski, G. 2004, *AcA*, 54, 153
- Prša, A. & Zwitter, T. 2005, *ApJ*, 628, 426
- Smith, J. A., Tucker, D. L., Kent, S., et al. 2002, *AJ*, 123, 2121
- Straizys, V. & Kuriliene, G. 1981, *Ap&SS*, 80, 353
- Talens, G. J. J., Spronck, J. F. P., Lesage, A.-L., et al. 2017, *A&A*, 601, 11
- van Hamme, W. 1993, *AJ*, 106, 2096
- Wilson, R. E. & Devinney, E. J. 1971, *ApJ*, 166, 605
- Zahn, J.-P. 1977, *A&A*, 57, 383

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