



Brazilian Journal of Physics

ISSN: 0103-9733

luizno.bjp@gmail.com

Sociedade Brasileira de Física

Brasil

Velásquez-Toribio, A. M.

Stability in a Class of Quartessence Models

Brazilian Journal of Physics, vol. 36, núm. 3A, septiembre, 2006, pp. 772-776

Sociedade Brasileira de Física

São Paulo, Brasil

Available in: <http://www.redalyc.org/articulo.oa?id=46413654034>

- How to cite
- Complete issue
- More information about this article
- Journal's homepage in redalyc.org

redalyc.org

Scientific Information System

Network of Scientific Journals from Latin America, the Caribbean, Spain and Portugal

Non-profit academic project, developed under the open access initiative

Stability in a Class of Quartessence Models

A. M. Velásquez-Toribio

Federal University of Rio de Janeiro (UFRJ), Institute of Physics,
P. O. Box 68528, CEP 21941-972, Rio de Janeiro, Brazil

Received on 13 April, 2006

Quartessence cosmological models with exponential and logarithmic equation of state are investigated using dynamical systems methods. We focus our analysis on the stability of these models.

Keywords: Cosmology; Quartessence; Dynamical systems; Stability

Progress in observational cosmology in the past few years suggests that the universe is dominated by two unknown components, namely, dark matter and dark energy. These two components have properties that are quite different from all the ordinary matter we know do exist in the physical world. Dark matter is a pressureless component, very important for structure formation, while dark energy, a negative pressure component, is responsible for the cosmic acceleration. From the point of view of simplicity, it would be interesting to describe both components with a single equation of state (EOS). This idea is the main motivation for quartessence models that try to unify dark matter and dark energy in a single component [1]. The most studied quartessence model is the Chaplygin gas model [2], [3]. Nevertheless other EOS for quartessence exist as, for instance, those suggested in the work by Reis et al. [4].

In this work, we employ the methods of qualitative theory of dynamical system [5], [6] to study quartessence models with exponential and logarithmic EOS. Recently, M. Szydlowski and W. Czaja [7] investigated the structural stability of FRW cosmology with generalized Chaplygin gas. Here, we extend their analysis to other quartessence cosmologies.

This paper is organized as follows. In Section II, we briefly describe the dynamical systems techniques we use to analyze quartessence cosmological models. In Section III we analyze the critical points of the considered quartessence models. In Section IV we investigate the stability of the system using the Lyapunov's method. Finally, in Section V we discuss the results.

I. STUDY OF QUARTESSENCE MODELS WITH DYNAMICAL SYSTEMS

The main goal of quartessence cosmological models is to describe with a single equation of state, dark matter and dark energy. The prototype quartessence model is the generalized Chaplygin Gas (GCG) [3]. It is easy to show that for very high redshift these models behave like pressureless dark matter, while in the opposite limit the models asymptotically tend to a cosmological constant. Different aspects of the GCG models have been analyzed in the literature [2], [8], [9]. Recently, Reis et al. [4] suggested two other equations of state with the same asymptotic limits, namely Exponential and Logarithmic Quartessence [10]. It is clear that, in principle, many other equations of state can be used, but for the sake of

simplicity, and because they capture the main properties of this class of models, here we limit our analysis to these two cases. The considered models have the following EOS:

Exponential Quartessence:

$$P = -M^4 e^{-\frac{\rho\alpha}{M^4}}, \quad (1)$$

Logarithmic Quartessence:

$$P = -\frac{M^4}{\left[\ln\left(\frac{\rho}{M^4}\right)\right]^\alpha}, \quad (2)$$

where ρ is the density energy, P the pressure, M has dimension of mass and α is a dimensionless parameter.

We denote by H the Hubble parameter and the quantities (H, ρ) are chosen as the phase space variables. We consider a homogeneous and isotropic Friedmann-Robertson-Walker (FRW) metric and we define our two-dimensional dynamical system with the following equations:

$$\dot{H} = -H^2 - \frac{1}{6}(\rho + 3P), \quad (3)$$

$$\dot{\rho} = -3H(\rho + P), \quad (4)$$

where the dot denotes time derivative. We split in two parts, our discussion of the character of the critical points:

A) static case ($H=0$)

and

B) non-static case ($H \neq 0$).

We consider non-degenerate critical points where the behavior of the system in the neighborhood of the critical point is qualitatively equivalent to the behavior of its linear part. The linearization matrix of the system at the critical point (H_0, ρ_0) is:

$$A = \begin{pmatrix} -2H & -\frac{1}{6}\frac{d}{d\rho}(\rho + 3P) \\ -3(\rho + P) & -3H\frac{d}{d\rho}(\rho + P) \end{pmatrix}_{(H_0, \rho_0)}$$

As usual, the characteristic equation is $\det(A - \lambda I) = 0$, where I is the identity matrix. The trace and the determinant for static critical points are [7],

$$\text{Tr}A = 0 \quad (5)$$

$$\det A = -\frac{1}{2}(\rho + P)\frac{d}{d\rho}(\rho + 3P)|_{(H_0, \rho_0)}. \quad (6)$$

For non-static critical points we obtain:

$$TrA = -H(2 + 3\frac{d}{d\rho}(\rho + P))|_{(H_0, \rho_o)} \quad (7)$$

$$detA = 6H^2\frac{d}{d\rho}(\rho + P)|_{(H_0, \rho_o)} \quad (8)$$

Here $|_{(H_0, \rho_o)}$ denotes the value of the equation at the critical point.

Since we are working with a two-dimensional dynamical system, the stability of the solution can be analyzed based on the sign of the trace and the discriminant [12]:

1. If $detA < 0$, then the eigenvalues are real, have opposite signs and the critical point is a saddle point.
2. If $detA > 0$ and the discriminant: $D = (TrA)^2 - 4detA > 0$, then the eigenvalues are real, have the same sign and the critical point is a node. If $TrA > 0$ the critical point is an unstable point(node). If $TrA < 0$ the critical point is stable.
3. If $detA > 0$ and the discriminant: $D = (TrA)^2 - 4detA < 0$, then the eigenvalues are complex conjugates and the critical point is a focus. If $TrA > 0$ the critical point is an unstable focus. If $TrA < 0$ it is a stable focus. In the case in which the eigenvalues are purely imaginary, the critical point is a stable neutral center.

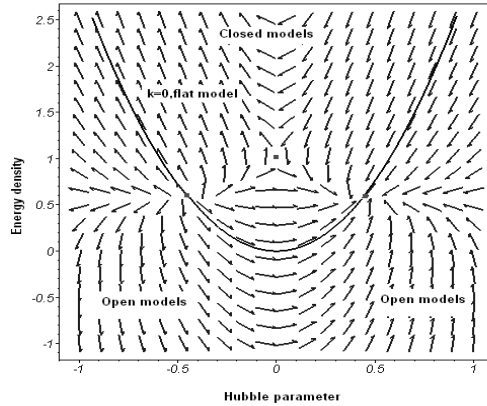


FIG. 1: The vector field portrait corresponding to the exponential quartessence case. For the figure we used $\alpha = M^4 = 1$

II. TWO QUARTESSENCE MODELS

The two cosmological models that we are interested in have the property that their equation of state $w = \frac{p}{\rho}$ is convex, i.e. $\frac{d^2 p}{d\rho^2} < 0$. In this case, the maximum value of the sound speed ($c_s^2 = \frac{dp}{d\rho}$) is reached at the minimum value of the density. Some of the properties of these models have been discussed in [10]. In this section we consider the stability with the methods of qualitative analysis described in Section II.

A. Exponential Quartessence

For quartessence models, at early times, when ρ is large, we need $w \approx 0$, therefore, from Eq. (1), the parameter α can not be negative. The minimal energy density of this model is given by $\rho_{min} = M^4(\frac{W(\alpha)}{\alpha})$, where $W(x)$ is the Lambert function defined to be the multi-valued inverse of the function $W \rightarrow We^W$ [10, 11]. Thus, when $\rho = \rho_{min}$, the equation of state parameter assumes the value $w = -1$ (cosmological constant). The standard Λ CDM is recovered when we substitute $\alpha = 0$ in Eq. (1).

In the considered model the singular static point is given by:

$$H_o = 0, \quad \rho_o = M^4 \frac{W(3\alpha)}{\alpha} \quad (9)$$

After substituting Eqs. (1) and (9) into (6), we obtain:

$$detA = \frac{-1}{2} \left(\frac{M^4}{\alpha} W(3\alpha) - M^4 e^{-W(3\alpha)} \right) (1 + 3\alpha e^{-W(3\alpha)}) \quad (10)$$

As the $TrA=0$, then, if $detA > 0$, then the eigenvalues of the critical points have complex values, that correspond to the stable neutrally center type. Otherwise, as discussed above, if the determinant is negative the critical point represent a saddle point.

Let us now analyze the non-static critical points: $H \neq 0$. Equaling equations (3) and (4) to zero, we determine the non-static critical points :

$$\rho_o = \frac{M^4}{\alpha} W(\alpha), \quad H_o^2 = -\frac{1}{6} \left(\frac{M^4}{\alpha} W(\alpha) - 3M^4 e^{-W(\alpha)} \right), \quad (11)$$

Using Eqs. (7) and (8), we also obtain:

$$detA = 6H_o^2(1 + \alpha e^{-W(\alpha)}) \quad (12)$$

and

$$TrA = -H_o(2 + 3(1 - \alpha e^{-W(\alpha)})) \quad (13)$$

It follows from Eq. (12) that the determinant is always positive. In the case of an expanding universe, ($H > 0$), the trace is negative. Therefore, the non-static critical point (11) behaves as an asymptotically stable node if the discriminant is positive ($D > 0$) or as an asymptotically stable focus if $D < 0$. In a contracting Universe, such that the Hubble parameter is negative, the critical points are unstable. In Fig. 1 we show the vector field portrait of the dynamical system formed by equations (3) and (4) for the exponential quartessence EOS. A striking features of Fig.1 is that the static critical point is associated with a saddle point (from equation (10) the $detA$ is negative for $\alpha \in [0, 1]$) and represents an Einstein static universe, whereas the non-static critical points are associated with a de Sitter(deS) universe. In the Figure 1 the trajectory of the flat model ($k=0$) is given by the parabola . The two deS nodes are attractors (one stable and the other one unstable) and the phase portrait is symmetric when H changes its sign.

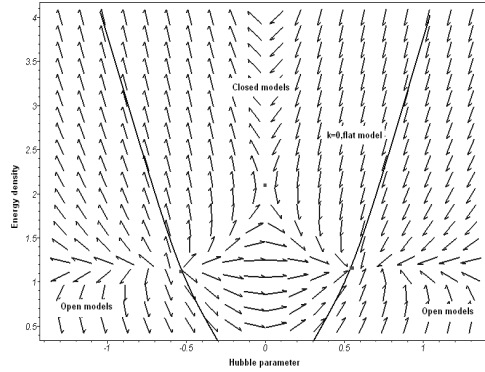


FIG. 2: The vector field portrait corresponding to the logarithmic quartessence case. For the figure we used $\alpha = 0.5, M^4 = 1$.

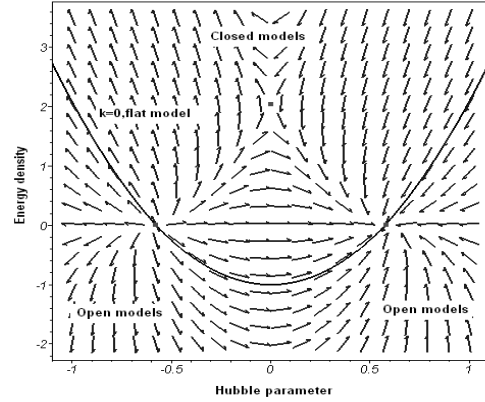


FIG. 3: The vector field portrait corresponding to the FRW model with cosmological constant.

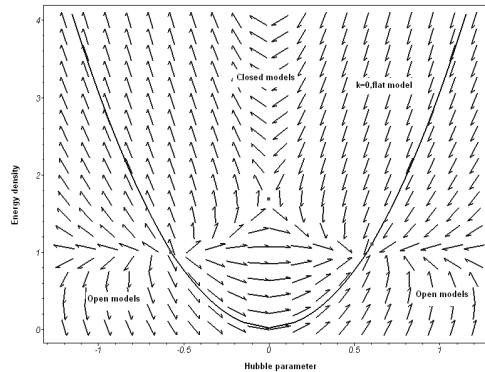


FIG. 4: The vector field portrait corresponding to the FRW model with Chaplygin gas. For the figure we used $\alpha = M^4 = 1$.

B. Logarithmic Quartessence

The logarithmic quartessence model EOS is given by equation (2). In the following we consider only the non-static case. The minimum energy density is given by $\rho_{min} = M^4 e^{\alpha W(\alpha^{-1})}$ and the adiabatic sound speed is $c_{smax}^2 = 1/W(\alpha^{-1})$. Note that again Λ CDM is recovered if we substitute $\alpha = 0$ in equation (2) [10]. Equalizing to zero Eqs. (3) and (4) and by using the equation of state (2) we obtain the following critical point (non-static case):

$$\rho_o = M^4 e^{\alpha W(\alpha^{-1})}, \quad H_o^2 = -M^4 \frac{1}{6} [e^{\alpha W(\alpha^{-1})} - 3(\alpha W(\alpha^{-1})^\alpha)] \quad (14)$$

The value of the trace and the determinant are:

$$detA = 6H_o^2 \left\{ 1 + \frac{\alpha M^4}{(\alpha W(\alpha^{-1})^{\alpha+1})} e^{-\alpha W(\alpha^{-1})} \right\} \quad (15)$$

$$TrA = -H_o \left\{ 2 + 3 \left(1 + \frac{\alpha M^4}{\alpha W(\alpha^{-1})^{1+\alpha}} e^{-\alpha W(\alpha^{-1})} \right) \right\} \quad (16)$$

Similarly, as we did in the exponential case, we deduce from equation (15) that the determinant is always positive and, from the Eq. (16) the trace is negative for an expanding universe ($H > 0$). Therefore, the critical points (14) behave like an asymptotically stable node if $D > 0$ or as an asymptotically stable focus if $D < 0$. In Figure 2 we show the phase portrait in the logarithmic quartessence case.

In Figures 3 and 4 we show the phase portrait of the FRW models with a cosmological constant (positive) and Chaplygin gas respectively, let us note the equivalence of the phase portraits in the physical domains ($\rho \geq 0$) of the four cosmological models (Figs. 1, 2, 3 and 4) i.e. the structure of the phase space does not depend of the form of EOS. The behavior for the phase portraits are due to the fact that the stability of the critical point depends on the energy condition as discussed in [7].

III. LYAPUNOV FUNCTION AND QUARTESSENCE

There are different approaches to investigate the stability of dynamical systems. In the following we shall consider the Lyapunov function method.

The existence of a Lyapunov function, F , guarantees the stability of the dynamical system around x_0 . Furthermore, if the inequality $\frac{dF}{dt} \leq 0$ holds in a neighborhood U_{x_0} , then x_0 is asymptotically stable. For a mathematical discussion of this method see e.g. refs. [13–15].

For the two-dimensional dynamical system (3) and (4) we propose the following Lyapunov's function:

$$F(\rho, H) = \rho + \frac{H^2}{2} > 0 \quad (17)$$

The above function is positive definite if $\rho > 0$. From equations (3)-(4) we have critical points if $H = 0$ and $w = -\frac{1}{3}$ or

$H^2 = \frac{\rho}{3}$ with $w = -1$, and they represent asymptotic states of the universe.

We also assume an expanding Universe and, in order to assure the asymptotic stability of the system of equations, we require that F first total derivative be negative. Therefore, we can write:

$$\frac{dF}{dt} = \frac{d\rho}{dt} + H \frac{dH}{dt} \leq 0 \quad (18)$$

In an expanding universe ($H > 0$) with $w \geq -1$ the energy density obeys the condition $\frac{d\rho}{dt} \leq 0$. Using Eq. (3) the quantity $\frac{dH}{dt}$ can be cast in the following form

$$\frac{dH}{dt} = -\left[\frac{\rho}{3} - \frac{k}{a^2} + \frac{(\rho + 3P)}{6}\right] \quad (19)$$

On the other hand, if we now use the definitions of the density parameters: $\Omega_k = -k/a^2H^2$, $\Omega_Q = \rho/3H^2$ and since $\Omega_k + \Omega_Q = 1$, it is easy to show that:

$$\frac{dH}{dt} = -\frac{\rho}{3}\left[\frac{3}{2}(1+w) + \frac{(1-\Omega_Q)}{\Omega_Q}\right] \quad (20)$$

From Eqs. (20) and (4) we have,

$$\frac{dF}{dt} = -\left(\frac{7}{2}(1+w)H\rho + H\rho\frac{(1-\Omega_Q)}{\Omega_Q}\right) \quad (21)$$

$$\Omega_Q \geq -\frac{2}{7w+5} \quad (22)$$

The last equation is the condition for stability of the dynamical system (3)-(4) with quartessence EOS in the case $w \geq -1$. The important point here is the existence of the Lyapunov function satisfying condition (22). In summary, the function F has a minimum (when $\Omega_Q = -\frac{2}{7w+5}$) on the critical point and there is a neighborhood (the attraction basin) with $H > 0$ where F has non-positive time derivative. Hence, this asymptotic states of the universe are asymptotically stable.

IV. CONCLUSIONS

In the present paper we consider the stability of two quartessence models, one with an exponential EOS and the other with a logarithmic one. The dynamics is studied in the phase plane (H, ρ) at a finite domain, applying methods of the qualitative theory of dynamical systems. We considered hyperbolic critical points. For both, the exponential and logarithmic cases, if the Hubble function is positive the models are stable in the neighborhood of the critical points. Therefore, it is interesting that the requirement of stability for these quartessence models yields an Universe which accelerates, in accordance with observational evidences at present [17].

To study the asymptotic behavior of the above mentioned quartessence cosmological models we have applied the Lyapunov method. We have shown that exists a Liapunov function (F) given by Eq.(17). This function can be thought as a generalized energy function for the system, and it leads us to the stability condition given by Eq.(22).

In summary, in this work we showed that quartessence models with exponential and logarithmic EOS are stable in their physical domains. Therefore, since the real world is also dynamically stable [13], these quartessence models, from the theoretical point of view, in principle, might describe the late-time acceleration of the Universe. Finally we remark that, for adiabatic perturbations, the matter power spectrum for these quartessence models present instabilities and oscillations [18]. However, as shown in [10], this kind of problem can be solved by considering a special type of non-adiabatic perturbation. We believe that these quartessence models deserve to be further investigated.

V. ACKNOWLEDGMENTS

The author wish to thank Ioav Waga for several useful discussions that helped us to improve this work and the Brazilian research agency for financial support: CNPq.

-
- [1] M. Makler, S. Q. Oliveira, and I. Waga, Phys. Lett. B **555**, 1 (2003).
 - [2] A. Kamenshchik, U. Moschella, and V. Pasquier, Phys. Lett. B **511**, 265 (2001); N. Bilic, G. B. Tupper, R. D. Viollier, Phys. Lett. B **535**, 17 (2002).
 - [3] M. C. Bento, O. Bertolami, and A. A. Sen, Phys. Rev. D **66**, 043507 (2002).
 - [4] R. R. R. Reis, M. Makler, and I. Waga, Phys. Rev. D **69**, 101301 (2004).
 - [5] O. I. Bogoyavlensky, Methods in Qualitative Theory of Dynamical Systems in Astrophysics and Gas Dynamic (Springer - Verlag, Berlin 1985).
 - [6] J. Wainwright and G. F. R. Ellis, Dynamical System in Cosmology, Cambridge University Press (1997).
 - [7] M. Szydlowski, W. Czaja, Phys. Rev. D **69**, 023506 (2004).
 - [8] M. Makler, S. Q. Oliveira, and I. Waga, Phys. Rev. D **64**, 123521 (2003).
 - [9] P. P. Avelino, L. M. G. Beca, J. P. M. de Carvalho, C. J. A. P. Martins, and P. Pinto, Phys. Rev. D **67**, 023511 (2003).
 - [10] R. R. R. Reis, M. Makler, and I. Waga, Class. Quantum. Grav **22**, 353 (2005).
 - [11] S. R. Cranmer, American J. Phys., **72**, 1397 (2004).
 - [12] D. V. Anosov and V. I. Arnold, *Dynamical Systems I*, Encyclopaedia of Mathematical Sciences, Volume 1 (1980).
 - [13] L. H. A. Monteiro, *Sistemas Dinamicos*, Editora Livraria da Fisica (2002).
 - [14] L. Perko, *Differential Equations and Dynamical Systems*, Second Edition, Springer-Verlag (1996).

- [15] T. C. Charters, A. Nunes, and J. P. Mimoso, *Class. Quant. Grav.* **18**, 1703 (2001).
- [16] D. J. Albers, J. C. Sprott, DOI: SFI-WP 04-02-007, SFI Working Papers (2003).
- [17] S. Perlmutter *et al.*, *Nature* **319**, 51 (1998); S. Perlmutter *et al.*, *Astrophys. J.* **517**, 565 (1999); A. G. Riess *et al.*, *Astron. J.* **116**, 1009 (1998).
- [18] H. Sandvik, M. Tegmark, M. Zaldarriaga, and I. Waga, *Phys. Rev. D* **69**, 123524 (2004).